Testing a quantitative model for dark energy

Peter L. Biermann^{1,2,3,4,5} & Benjamin C. Harms³

¹ Dept. of Phys. & Astr., Univ. of Alabama, Tuscaloosa, AL, USA, and MPI for Radioastronomy, Bonn, Germany; ³ Dept. of Phys., Karlsruher Institut für Technologie KIT, Germany, ⁴ Dept. of Phys., Univ. of Alabama at Huntsville, AL, USA; ⁵ Dept. of Phys. & Astron., Univ. of Bonn, Germany;

plbiermann@mpifr-bonn.mpg.de, bharms@bama.ua.edu



Figure 1 Integral mass function corrected for Hubble type sampling, 2928 objects, the slope of the lines is: red -2.0 fitting > $10^8 M_{\odot}$, and blue -1.0 fitting between $10^7 M_{\odot}$ and $10^8 M_{\odot}$. See Caramete & PLB, 2010 AA. This mass function suggests that black holes start near 3 $10^6 M_{\odot}$, possibly at redshift of order $\lesssim 50$, and grow by merging (see PLB & Kusenko 2006 PRL). Lower mass objects included here are probably nuclear star clusters.

Bekenstein and consequences?

Bekenstein (1973 PRD) writes: Black hole entropy must be a property of the universe. Entropy is information, and information needs a carrier. Gravitational waves?

$$\frac{S}{k_B} = N_{GW} = 4\pi \left(\frac{M_{BH}}{m_{Pl}}\right)^2 , \text{ zero spin}$$

ಲು

With
$$E_{GW} = \frac{1}{8\pi} \frac{\hbar c^3}{G_N M_{BH}} = \frac{c^2}{8\pi} \frac{m_{Pl}^2}{M_{BH}}$$

 $N_{GW} E_{GW} = \frac{1}{2} M_{BH} c^2$

Can black holes match dark energy?

- $N_{BH,0} = 10^{-2.2} \text{ Mpc}^{-3}$ to 1 Mpc⁻³ Systematic errors, statistical errors, merging dominant?
- $M_{BH} = 10^6 \,\mathrm{M_{\odot}}$ to $10^7 \,\mathrm{M_{\odot}}$ From cut in mass function, instability, and Galactic Center black hole (Caramete & PLB 2010 AA)
- $z_{\star} = 25$ to 70 (PLB & Kusenko 2006 PRL) Very early massive star formation? Black holes?

$$\frac{1}{2} N_{BH,0} M_{BH} c^2 (1+z_{\star})^3 \simeq 10^{-8.1} \,\mathrm{erg} \,\mathrm{cm}^{-3} = \mathrm{DE}$$
From 2 10⁶ M with γ 50 to 1 10⁷ M with

From $3 \cdot 10^{\circ} M_{\odot}$ with $z_{\star} = 50$ to $1 \cdot 10^{\circ} M_{\odot}$ with $z_{\star} = 30$ at the same density 1 Mpc^{-3}

Three questions - three answers?

Requirement:
$$E_{BH} = \frac{1}{2} M_{BH} c^2 \left(\frac{1+z_{\star}}{1+z}\right)^3$$

- Initial shell of gravitational waves, thought of as coherent burst (soliton).
- Question: Gravitons as a soliton shell? What frame of reference possible?
- It grows from itself, so like stimulated emission.

сл

- Question: What is the phase space density in this case?
- The energy comes from a background. Inspired by Randall & Sundrum (1999), but variable in time
- Question: What exactly is the nature of the background? Energy density?

Boltzmann equation derivation

The distribution function $\mathcal{N}(k, t)$ for coherent production of gravitons in a Friedman-Robertson-Walker (FRW) universe, frequency of the wave at emission ν_0

$$\frac{\partial}{\partial t} \mathcal{N}(\nu_0, t) = \kappa_0 R(t) \frac{c \, k_B \, T_{g0}}{h \, \nu_0} \frac{\partial}{\partial \nu_0} \mathcal{N}(\nu_0, t) ,$$

$$\kappa_0 = \frac{1}{2\pi} \int \frac{d^3 p}{(2 \, \pi)^3} \frac{|M|^2}{E(p)^2}$$
Adopt $|M|^2 \simeq \frac{energy}{mass} \left(\frac{1+z}{1+z_\star}\right) \frac{1}{R^3} \frac{1}{\tau_H}$
where R^{-3} for quadrupoles, $R = \left(\frac{1+z_\star}{1+z}\right)$. With conditions for coherence and some numerical simplifications

$$\frac{1}{2}M_{BH}c^2\left(\frac{1+z_\star}{1+z}\right)^3$$

Heuristic derivation

Four steps: 1) Reference frame $\Gamma = \frac{1}{2} \frac{r(z, z_{\star})}{l_{Pl}} \{H(z)\tau_{Pl}\}^{1/2},$ 2) occupation number first bracket $\left(N_{GW} \frac{\sigma_{Pl}}{4\pi R_s^2}\Gamma\right),$ 3) coherence patch $\pi\lambda_{BH}^2,$

4) one per coherence patch per Planck time in comoving frame (interaction probability unity with resonant 3D momentum phase space).

For stimulated emission familiar non-linear form (Kompaneets 1957):

$$\frac{dE}{d\tau} = \left(N_{GW}\frac{\sigma_{Pl}}{4\pi R_s^2}\Gamma\right)^2 \left(\frac{4\pi R_s^2}{\pi \lambda_{BH}^2}\Gamma\right) \frac{E_{GW}\Gamma}{\tau_{Pl}}$$

In observer frame we obtain again

$$\frac{1}{2}M_{BH}c^2\left(\frac{1+z_\star}{1+z}\right)^3$$

Summary

- Dark energy drives universe apart with acceleration.
- We propose: Dark energy is ensemble of coherent bursts (soliton-like) gravitational waves originally produced when first generation of super-massive back holes was formed.
- These solitons get their initial energy as well as keep up their energy density throughout evolution of universe by stimulating emission from background.

 ∞

- Key tests involve pulsar timing, clock jitter, the radio, X-ray, γ -ray and neutrino background, early formation of pure-disk galaxies, and magnetization of IGM.
- First steps are in (PLB & Harms) arXiv:1205.4016, and in arXiv:1302.0040.

Abstract

Dark energy drives the universe apart with acceleration. We propose to develop quantitative tests of a concept for dark energy, and to evolve it further to a full theory. The concept is that dark energy is the ensemble of coherent bursts of (soliton-like) gravitational waves originally produced when the first generation of super-massive back holes were formed. These solitons get their initial energy as well as keep up their energy density throughout the evolution of the universe by stimulating emission from a background. Key tests involve pulsar timing, clock jitter, the radio-, X-ray, gamma- and neutrino-background, and the early formation of pure-disk galaxies. First steps are in arXiv:1205.4016, and arXiv:1302.0040 (PLB & Harms).

Acknowledgements

Intense discussions with Laurențiu I. Caramete, László Á. Gergely, Gopal-Krishna, Pankaj Joshi, Biman B. Nath, and many others are gratefully acknowledged.

References

- [1] Appenzeller, I., & Fricke, K., Astron. & Astroph.
 18, 10 (1972a).
- [2] Appenzeller, I., & Fricke, K., Astron. & Astroph.
 21, 285 (1972b).
- [3] Bekenstein, J. D., *Phys. Rev.* **D** 7, 2333 (1973).
- [4] Biermann, P.L., Kusenko, A., Phys. Rev. Letters 96, 091301 (2006); astro-ph/0601004.

- [5] A graviton statistics approach to dark energy, inflation and black holes, Biermann, P. L., & Harms, B. C., eprint arXiv:1205.4016
- [6] A Comprehensive Model of Dark Energy, Inflation AND Black Holes Authors: Peter L. Biermann and Benjamin C. Harms, Proc. 13th Marcel Grossmann Meeting, Stockholm July 2012, eprint arXiv:1302.0040
- [7] The mass function of nearby black hole candidates, Laurentiu I. Caramete, Peter L. Biermann, Astron.
 & Astroph. 521, id.A55 (2010); arXiv:0908.2764
- [8] Eckart, A., et al. 2005 BH in GC mass
- [9] Frieman, J. A., Turner, M. S., & Huterer, D., Annual Rev. of Astron. & Astrophys. 46, 385 (2008).
- [10] Greene, J.E., Barth, A.J., & Ho, L.C., New Astron.
 Rev. 50, 739 (2006); arXiv:astro-ph/0511810.

[11] Heger, A., Woosley, S. E., & Baraffe, I., in Proc. The Fate of the Most Massive Stars, ASP Conf. Ser. Vol. 332, Eds. R. Humphreys & K. Stanek, p. 347 (2005)

- [12] Seven-year Wilkinson Microwave Anisotropy Probe (WMAP) Observations: Sky Maps, Systematic Errors, and Basic Results, Jarosik, N., et al. Astrophys. J. Suppl. 192, id. 14 (2011); arXiv:1001.4744
- [13] Randall, L., Sundrum, R., Phys. Rev. Letters 83, 3370 (1999a)

- [14] Randall, L., Sundrum, R., Phys. Rev. Letters 83, 4690 (1999b)
- [15] Sanders, R. H., Astrophys. J. 162, 791 (1970).
- [16] Spitzer, L., Jr., Astrophys. J. Letters 158, L139 (1969).